On the Symmetries of Extended X-Ray Emission from Planetary Nebulae

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ON THE ASYMMETRIES OF EXTENDED X-RAY EMISSION FROM PLANETARY NEBULAE

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ABSTRACT

Chandra X-Ray Observatory (CXO) images have revealed that the X-ray–emitting regions of the molecule-rich young planetary nebulae (PNe) BD +30° 3639 and NGC 7027 are much more asymmetric than their optical nebulosities. To evaluate the potential origins of these X-ray asymmetries, we analyze X-ray images of BD +30° 3639, NGC 7027, and another PN resolved by CXO, NGC 6543, within specific energy bands. Image resolution has been optimized by subpixel repositioning of individual X-ray events. The resulting subarcsecond resolution images reveal that the soft ($E < 0.7$ keV) X-ray emission from BD +30° 3639 is more uniform than the harder emission, which is largely confined to the eastern rim of the optical nebula. In contrast, soft X-rays from NGC 7027 are highly localized and this PN is more axially symmetric in harder emission. The broadband X-ray morphologies of BD +30° 3639 and NGC 7027 are highly anticorrelated with their distributions of visual extinction, as determined from high-resolution, space- and ground-based optical and infrared imaging. Hence, it is likely that the observed X-ray asymmetries of these nebulae are due in large part to the effects of nonuniform intranebular extinction. However, the energy-dependent X-ray structures in both nebulae and NGC 6543—which is by far the least dusty and molecule-rich of the three PNe and displays very uniform intranebular extinction—suggest that other mechanisms, such as the action of collimated outflows and heat conduction, are also important in determining the detailed X-ray morphologies of young PNe.

Subject headings: planetary nebulae: individual (BD +30° 3639, NGC 7027) — stars: AGB and post-AGB — stars: mass loss — stars: winds, outflows — X-rays: ISM

On-line material: color figures

1. INTRODUCTION

Models of the formation of planetary nebulae (PNe) have long predicted that these objects, representing very late stages in the deaths of intermediate-mass ($1\text{–}8\,M_\odot$) stars, should emit X-rays. Such emission should arise in shocks at the interface between an active wind from the PN core (or its companion) and material ejected when the progenitor was on the asymptotic giant branch (AGB). Thus, extended X-ray emission, if present, likely traces the very processes responsible for sculpting PNe (for recent discussions of PN-shaping mechanisms, see, e.g., Frank 1999; Gardiner & Frank 2001; Kastner, Soker, & Rappaport 2000a; Soker & Rappaport 2000).

The Chandra X-Ray Observatory (CXO), with its unprecedented spatial resolution, has now provided the first conclusive evidence of such extended X-ray emission from nebular gas, in the form of striking X-ray imagery of the young PNe BD +30° 3639 (Kastner et al. 2000b, hereafter KSVD00), NGC 7027 (Kastner, Vrtilek, & Soker 2001, hereafter KVS01), and NGC 6543 (Chu et al. 2001). Observations of the PN NGC 7009 by *XMM-Newton*, which combines high sensitivity with good spatial and spectral resolution, also reveal marginally extended X-ray-emitting gas (Guerrero, Chu, & Gruendl 2002). The PN NGC 7293 (the Helix), known to exhibit relatively hard X-ray emission, does not display evidence for extended X-ray emission in CXO imaging (Guerrero et al. 2001); instead, this PN and NGC 6543 contain pointlike X-ray sources with temperatures of a few $\times 10^6$ K, possibly due to magnetic activity on companions to their central stars (Guerrero et al. 2001; Gruendl et al. 2001; Soker & Kastner 2002).

While revealing the diffuse nature of the X-ray emission from some PNe, these first CXO and *XMM-Newton* observations of PNe are notable and surprising in several respects. Of particular interest is the result that the X-ray morphologies of the young PNe BD +30° 3639 and NGC 7027 are decidedly asymmetric, much more so than their optical nebulosities. Although the extended emission detected in the CXO images of these PNe (and NGC 6543 and 7009) underscores the importance of strong shocks in shaping planetaries, the asymmetric structures observed by CXO in BD +30° 3639 and NGC 7027 cannot be easily explained in terms of “fossil,” spherical AGB envelopes acted on by isotropic white dwarf winds. However, as BD +30° 3639 and NGC 7027 are also among the youngest, dustiest, and most molecule-rich of known PNe, it is important to assess the extent to which their observed X-ray asymmetries represent intrinsic plasma density or temperature structure inhomogeneities rather than, e.g., the effects of intervening X-ray absorption.

In this paper, we explore the origin of asymmetries of extended X-ray emission from PNe (in a companion paper, Soker & Kastner 2003, we examine models to explain the observed X-ray luminosities and temperatures of PNe). In
particular, we consider the role of nonuniform intranebular extinction in determining the X-ray morphologies of BD +30°3639, NGC 7027, and NGC 6543. In doing so, we take advantage of the subarcsecond imaging potential of CXO and the energy resolution of CXO’s Advanced CCD Imaging Spectrometer (ACIS). In § 2, we present archival optical and near-infrared images of BD +30°3639 and NGC 7027 that are useful in deducing their distributions of intranebular extinction, and we summarize the application of a subpixel image restoration technique to CXO images of these nebulae and NGC 6543. Results, including presentation of superresolved energy-band images of all three PNe and comparisons of their X-ray surface brightnesses and spatial distribution of intranebular extinction (as inferred from the optical/near-IR data), are presented in § 3. In § 4, we discuss our main findings, and § 5 contains a summary.

2. OBSERVATIONS AND DATA PROCESSING

2.1. Optical and Infrared

Images of BD +30°3639 in the transitions of Hα (0.6563 μm) and Pα (1.87 μm) utilized in the present analysis (§ 3) were obtained with the Hubble Space Telescope (HST) using the Wide Field Planetary Camera 2 (WFPC2) and Near-Infrared Camera and Multi-Object Spectrometer (NICMOS), respectively. These images appear in Figure 1. The Hα and Pα images were first presented in Sahai & Trauger (1998) and Latter et al. (2000a), respectively; in addition, HST narrowband images covering many diagnostics emission lines were presented in Arnaud, Borkowski, & Harrington (1996) and Harrington et al. (1997). The reader is referred to these papers for details concerning the structure and physical conditions of BD +30°3639 as ascertained from HST imaging. As is readily apparent from Figure 1, however, the inner nebula consists of a bright elliptical shell with major and minor axes of ~4″ and ~3″, respectively (this inner nebula is surrounded by a much larger, fainter Hα halo; Sahai & Trauger 1998). There is a strong gradient in extinction across this region of the nebula, with many conspicuous knots and clumps of denser gas apparent in projection against the bright elliptical shell (Harrington et al. 1997).

Images of NGC 7027 in the transitions of Brγ (2.16 μm) and Brα (4.05 μm) utilized in the present analysis (§ 3) were obtained with the National Optical Astronomy Observatory\(^2\) 4 m telescope and Cryogenic Optical Bench (COB) at Kitt Peak National Observatory in 1995 September (for details concerning COB imaging on the 4 m, see Weintraub et al. 1996). These images appear in Figure 2. Like BD +30°3639, NGC 7027 displays a bright elliptical shell, with a fainter surrounding halo; the major and minor axes of the shell are ~10″ and ~7″, respectively. While extensive near-infrared imagery of NGC 7027 has been obtained in the 2–3 μm wavelength range (e.g., Graham et al. 1993; Kastner et al. 1994, 1996; Latter et al. 2000b; Cox et al. 2002), the Brα image presented here represents the longest wavelength image of the nebula presently available at ~1″ resolution. As we demonstrate in § 4, this image is valuable for examining, at high spatial resolution, highly obscured regions of the nebula.

2.2. X-Ray

BD +30°3639 and NGC 7027 were observed by CXO in 2000 March and June, with net integration times of 18.8 and 18.2 ks, respectively. Both PNe were imaged with the central back-illuminated CCD in the ACIS array (ACIS-S3). These observations were first presented in KSVD00 and KVS01. NGC 6543 was observed with ACIS-S3 for 46.0 ks in 2000 May (Chu et al. 2001). In this paper, we make use of these CXO/ACIS-S3 event data as reprocessed by the Chandra X-ray Center (CXC) in 2000 December for BD +30°3639 and in 2001 January for NGC 7027 and NGC 6543 (Figs. 3, 4, and 5, respectively).

2.2.1. Event Position Refinement

The FWHM of the core of the point-spread function (PSF) of the CXO High-Resolution Mirror Assembly (HRMA) is nearly equivalent to the ACIS-S3 pixel size, 0″.49.\(^6\) Because (1) the telescope position is dithered across

\(^5\) The National Optical Astronomy Observatory is operated by the Association of Universities for Research in Astronomy (AURA), Inc., under cooperative agreement with the National Science Foundation.

Fig. 2.—Brγ 2.16 μm (left) and Brα 4.05 μm (right) images of NGC 7027, obtained with the COB on the Kitt Peak 4 m telescope.

Fig. 3.—Left: CXO X-ray image of BD +30° 3639, obtained by binning events before removing position randomization or applying subpixel event position corrections. Center: Image of BD +30° 3639 obtained by binning events after removing event position randomization. Right: Image of BD +30° 3639 obtained by binning events after removing randomization and applying subpixel event position corrections. Images here and in Figs. 4 and 5 are presented in a log scale, with the bars representing the gray-scale mapping to log intensity.

Fig. 4.—Left: CXO X-ray image of NGC 7027, obtained by binning events before removing position randomization or applying subpixel event position corrections. Center: Image of NGC 7027 obtained by binning events after removing event position randomization. Right: Image of NGC 7027 obtained by binning events after removing randomization and applying subpixel event position corrections.
the source during an observation and is known to an accuracy of better than $0^\prime\prime.3$, (2) the charge cloud size generated by an incident photon is much smaller than the ACIS pixel size (Tsunemi et al. 2001), and (3) the ACIS flight software records the distribution of charge among pixels for each candidate X-ray event (encoded as “FLTGRADE,” the bitmap value of pixels above the “split threshold”). It is possible (in principle) to reconstruct the position of each detected X-ray to subpixel accuracy (e.g., Tsunemi et al. 2001). We have thus applied event position corrections to the reprocessed data obtained for BD +30\degree C143639, NGC 7027, and NGC 6543, following the general method outlined in Tsunemi et al., as implemented, refined, and tested by Li et al. (2002). That is, we move the position of a photon from the default (pixel center) to its inferred landing location in detector coordinates, according to the event grade (see below). We then use the telescope pointing and spacecraft roll angle history for the observation to project the revised event locations to sky coordinates.

The Tsunemi et al. (2001) method relies on rejection of all events except those with charges distributed among 3 or 4 pixels; these “corner split” events have better subpixel accuracy in both directions (horizontal and vertical) than single events and in one direction than 2 pixel events, thereby maximizing the potential improvement in spatial resolution. However, corner splits constitute only about 10% of the total number of events for a typical ACIS-S3 observation of a soft source (e.g., Table 1). Thus, to boost the signal-to-noise ratio, Li et al. (2002) include single-pixel as well as 2 pixel events for which the charge is split either horizontally or vertically. As we demonstrate below, the improvement in image quality remains significant for this more inclusive FLTGRADE selection. In calculating new photon landing positions in detector coordinates, Li et al. assume that single-pixel events correspond to photons absorbed at the event pixel center, 2 pixel vertical or horizontal split event photons land at the centers of the split boundaries, and the corner split event photons land at the split corners. Before applying this event position relocation algorithm, position randomization within an ACIS pixel (performed as part of the standard CXC event processing pipeline) must be reversed.

In Table 1, we summarize the event FLTGRADE distributions for BD +30\degree 3639, NGC 7027, and NGC 6543. In compiling these statistics, we have selected those events with nominal energies less than 3.0 keV (as determined from calibrations produced by the standard CXC ACIS processing pipeline) that lie within the BD +30\degree 3639, NGC 7027, and NGC 6543 source regions as defined in KVS01, KSVD00, and Chu et al. 2001, respectively. It can be seen that our event selection criteria retain $\geq 95\%$ of events in the source regions (the selected events necessarily include a small percentage of background events).

3. RESULTS

3.1. Subpixel Image Reconstruction

Results of application of subpixel event relocations to the PN event data are displayed in Figures 3, 4, and 5. In each

![Image](image_url)

**Fig. 5.** Left: CXO X-ray image of NGC 6543, obtained by binning events before removing position randomization and applying subpixel event position corrections. Center: Image of NGC 6543 obtained by binning events after removing event position randomization. Right: Image of NGC 6543 obtained by binning events after removing position randomization and applying subpixel event position corrections.

### TABLE 1

**Event Grade Distributions for Planetary Nebulae Observed by CXO**

<table>
<thead>
<tr>
<th>Event Type</th>
<th>FLTGRADE</th>
<th>NGC 6543</th>
<th>BD +30\degree 3639</th>
<th>NGC 7027</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td>Number</td>
<td>%</td>
<td>Number</td>
</tr>
<tr>
<td>All.............</td>
<td>(All)</td>
<td>1755</td>
<td>...</td>
<td>4868</td>
</tr>
<tr>
<td>Single pixel...</td>
<td>0</td>
<td>830</td>
<td>47.3</td>
<td>1934</td>
</tr>
<tr>
<td>2 pixel.........</td>
<td>2, 8, 16, 64</td>
<td>764</td>
<td>43.5</td>
<td>2233</td>
</tr>
<tr>
<td>Corner split...</td>
<td>10, 11, 18, 22, 72, 80, 104, 208</td>
<td>127</td>
<td>7.2</td>
<td>456</td>
</tr>
<tr>
<td>Selected........</td>
<td>0, 2, 8, 10, 11, 16, 18, 22, 64, 72, 80, 104, 208</td>
<td>1721</td>
<td>98.1</td>
<td>4623</td>
</tr>
</tbody>
</table>
figure, the original image—i.e., the image obtained by spatially binning the Level 1 event data provided via standard event processing by the CXC, after filtering on energy and FLTGRADE (§2.2.1)—is presented in the left-hand panel. Images constructed from “unrandomized” event positions are displayed in the center panels of each figure; in both these unrandomized images and the “original” images constructed from Level 1 events, the pixel size is set to the intrinsic pixel size of ACIS, 0.049. Finally, the right-hand panels of each figure display images constructed after relocating events according to the subpixel event repositioning algorithm. These images were constructed for a pixel size of 0.025, which fully sampled the HRMA PSF at the energy range characteristic of the three PNe (0.3–3.0 keV).

The comparisons between “original,” “unrandomized,” and “event relocated” images in Figures 3, 4, and 5 illustrate the superior spatial resolution afforded by subpixel event repositioning. Perhaps the best example of such an improvement is the sharply decreased FWHM of the central point source in NGC 6543 (compare the left- and right-hand panels of Fig. 5): before correction, we measure a FWHM of 1.12′, and after correction, we measure a FWHM of 0.80′. Similarly, after event relocation is applied, the X-ray outline of BD +30°3639 and the bright region of emission at its eastern edge appear sharper (Fig. 3, right). The number of counts per pixel is quite small in the reconstructed image of NGC 7027, so some of the structure in this image is due to small number statistics. Nevertheless, features in the brighter regions of these images (e.g., the compact bright region located northwest of the position of the central star) appear to correspond to features seen in optical and near-infrared images of this PN (§3.2 and 3.3).

The absolute astrometry of the reprocessed BD +30°3639 and NGC 7027 data is improved over that available to KSVD00 and KVS01 because of refined aspect solutions. Hence, we can use the reconstructed images of these nebulae to reexamine the correspondence of their X-ray– and optical-emitting regions. For BD +30°3639, the relative alignment of the Po image and the broadband (0.3–3.0 keV) reconstructed image constructed from the reprocessed CXO event data agrees to within ~0.75″, based on comparison of the outlying CXO image contours with the elliptical shell seen in Po (Fig. 6). For NGC 7027, we have adjusted the position of the X-ray image by ~1″, to better align the positions of broadband peak X-ray surface brightness with the bright elliptical shell in the Brγ image (Fig. 7). The resulting overlays of reconstructed broadband X-ray images on infrared images (Figs. 6 and 7) indicate that (1) the central stars of both PNe are confirmed to lie very near the center of their X-ray nebulosities, and (2) unlike both NGC 7293 and NGC 6543 (Guerrero et al. 2001), neither BD +30°3639 nor NGC 7027 clearly contains an X-ray–bright central star (a possibility left open for NGC 7027 by the preliminary analysis of KVS01).

3.2. Energy-resolved Images

In Figures 8, 9, and 10, we present, respectively, broadband and energy-resolved images of BD +30°3639, NGC 7027, and NGC 6543. The energy-resolved images span the energy ranges $E = 0.7$ keV (hereafter soft band), $0.7 \text{ keV} < E \leq 1.2$ keV (medium band), and $1.2 \text{ keV} < E \leq 3.0$ keV (hard band). Like the broadband images, the energy-resolved images also have been reconstructed at subpixel resolution, with 0.25 pixels, via the algorithm described in the previous section. Our analysis of the subpixel event repositioning algorithm suggests that the results of the algorithm are robust, regardless of the count rate of the source, since the algorithm operates on one photon at a time (Li et al. 2002). Hence, the subpixel reconstructions are reliable, despite the small number of counts (and, hence, low signal-to-noise ratio) in certain energy-resolved images (e.g., hard-band images of BD +30°3639 and NGC 6543 and all images of NGC 7027).

The results demonstrate that the diffuse X-ray emission morphologies of all three nebulae are highly energy dependent within the 0.3–3.0 keV band. In all four images of BD +30°3639, it is apparent that the X-ray emission is brightest toward its eastern rim (Fig. 8); however, this asymmetry is much more profound in the medium- and hard-band images than in the soft-band image. In the soft-band image, the peak of the surface brightness of BD +30°3639 lies close to (although ~1″ east of) the position of the central star, while in the medium- and hard-band images, the nebula is much
Fig. 7.—Left: Reconstructed broadband CXO X-ray image of NGC 7027. Right: Contours of X-ray surface brightness overlaid on the Brα image of Fig. 2. Contour levels are at 4, 8, 16, and 24 counts arcsec$^{-2}$. In each panel, offset (0, 0) corresponds to the position of the central star to $\sim$0.72 (as determined from comparison of the Brα image with HST near-infrared imagery; Latter et al. 2000b).

Fig. 8.—Energy-resolved images of BD $+30^\circ$3639. Top left: Broadband image, obtained from all events with energies $\leq$3.0 keV. Top right: Soft-band image, obtained from events with energies $\leq$0.7 keV. Bottom left: Medium-band image, obtained from events with energies between 0.7 and 1.2 keV. Bottom right: Hard-band image, obtained from events with energies between 1.2 and 3.0 keV. Images here and in Figs. 9 and 10 are presented in a log scale, with the bars representing the gray-scale mapping to log intensity.
less centrally peaked. In these images, the (clumpy) emission from just inside the rims of the optical nebula appears generally stronger than the emission from the core region, and there is a compact bright region at the eastern rim.

In contrast to BD +30\degree3639, the soft-band emission from NGC 7027 is highly localized (Fig. 9). This emission is entirely confined to a region in the northwest of the nebula that appears as a “hole” in high-resolution optical images (see, e.g., Fig. 1 of KVS01). Emission from the X-ray-emitting “lobe” projected northwest of the central star—a direction corresponding to blueshifted atomic and molecular gas (Cox et al. 2002)—dominates the nebula in both the soft- and medium-band images. In the hard-band image, the emission is more balanced between the northwest (forward-facing) and southeast (rearward-facing) X-ray lobes, and the northwest X-ray lobe has a smaller opening angle than in the medium-band image. The outline of NGC 7027 in this hard-band image appears more or less axisymmetric or, perhaps, point-symmetric.

Figure 10 demonstrates that all diffuse X-ray emission detected from NGC 6543 emerges at energies less than 1.2 keV; only the central star is detected in the hard-band image. This is consistent with the spectral analysis in Chu et al. (2001) and Guerrero et al. (2001), which indicates that the diffuse emission has a characteristic temperature \(\sim 10^6\) K, while the unresolved central source is somewhat hotter, at \(\sim 2 \times 10^6\) K. In the medium-band image, the emission appears confined to the edges of the soft-band nebulosity, with a clump of brighter emission at the extreme northern edge of the X-ray-emitting region.

3.3. X-Ray Surface Brightness versus Optical/IR Extinction

Assuming case B recombination, the hydrogen emission-line ratios \(I_{H\alpha}/I_{P\alpha}\) and \(I_{Br\gamma}/I_{Br\alpha}\) are relatively insensitive to temperature in the regime of interest for planetary nebulae, \(T \sim 10^4\) K (e.g., Osterbrock 1989, Table 4.2). Therefore, we can assume that deviations of the observed line ratios from their “canonical” case B recombination values are due to extinction. Hence, assuming a standard interstellar dependence of extinction on wavelength (e.g., Osterbrock 1989, Table 7.2), we use the spatial distribution of the ratios \(I_{H\alpha}/I_{P\alpha}\) and \(I_{Br\gamma}/I_{Br\alpha}\), as derived from the ratios of images presented in § 2.2, to infer visual extinction (\(A_V\)) as a function of position within BD +30\degree3639 (Fig. 11) and NGC 7027 (Fig. 12), respectively.
Harrington et al. (1997) and Robberto et al. (1993) employed techniques similar to the preceding to infer the spatial distribution of $A_V$ for BD $+30\degree 3639$ and NGC 7027, respectively. Their results are qualitatively similar to those derived here. The extinction across both nebulae can be quite large, however, and use of near-infrared images provides a better “lever arm” for deducing $A_V$ in very highly obscured regions. For example, the Robberto et al. (1993) extinction map of NGC 7027, constructed from optical hydrogen recombination line imaging, is limited to regions with $A_V < 5$, whereas the $A_V$ map constructed from longer wavelength hydrogen transitions available to COB (Fig. 12) remains sensitive to $A_V \sim 15$.

The distributions of $A_V$ are similar for the two nebulae, in that the largest values of extinction lie near the nebular perimeters and extinction “holes” are observed toward the nebular interiors. Generally, however, we find $A_V$ within NGC 7027 (for which $A_V$ lies largely in the range $1.5 \leq A_V \leq 15$) to be a factor of 3–4 larger than $A_V$ within BD $+30\degree 3639$ (for which $0.5 \leq A_V \leq 5$). Also, the dark lane apparent in high-resolution optical images of NGC 7027 (e.g., KVS01, their Fig. 1) appears as an enhancement of $A_V$ in Fig. 12. This is consistent with the hypothesis (KVS01) that BD $+30\degree 3639$ and NGC 7027 share a common structure—bipolar, with a dense equatorial region—but that NGC 7027 is viewed at an intermediate inclination angle, whereas BD $+30\degree 3639$ is viewed nearly pole-on.

Notably, the X-ray surface brightnesses of both nebulae are strongly anticorrelated with the local value of $A_V$. In each case, the peak of X-ray emission lies very near the minimum in $A_V$, and little or no X-ray emission is observed toward regions of highest $A_V$. There is a gradient of $A_V$ in interior regions of BD $+30\degree 3639$ (i.e., regions of the nebula interior to the bright shell seen in H recombination lines), such that there is greater extinction toward the southwestern side of the central star (a result independently obtained by Harrington et al. 1997), and there is an extinction hole $1''-2''$ to the eastern side of the central star. The X-ray emission, correspondingly, is much brighter to the east than to the west of the star, and furthermore, the contours of X-ray surface brightness closely follow the $A_V$ distribution in detail (Fig. 11). In NGC 7027, the relationship between $A_V$ and X-ray surface brightness is also strong, with the lowest contours of X-ray emission very closely tracing the regions of large $A_V$ (Fig. 12). In particular, the “pinched waist” apparent in the X-ray emission morphology lies at the

Fig. 10.—Energy-resolved images of NGC 6543. Energy ranges for panels are as in Fig. 8.
position of the enhancement of $A_V$ that appears to define the equatorial plane of the nebula. Also, the strongest X-ray emission from NGC 7027 lies very near—although slightly farther from the central star than—the extinction hole apparent to the northwest of the star.

We have also constructed an $H\alpha/H\beta$ line ratio map of NGC 6543 (not shown), from image data available in the HST archive. In stark contrast to the line ratio maps of BD $+30\degr$3639 and NGC 7027, from which the $A_V$ maps in Figures 11 and 12 and were constructed, the $H\alpha$ line ratio map of NGC 6543 is smooth and featureless across its central X-ray-emitting region, indicating that any extinction of the X-ray nebula is quite spatially uniform. Furthermore, the measured $H\alpha/H\beta$ ratio across this region, $2.5 \pm 0.3$, agrees within uncertainty with the value expected for case B recombination, $\approx 2.8$ (Osterbrock 1989). This result is consistent with the very small inferred extinction toward NGC 6543 as measured by Tylenda et al. (1992) and with the observation that X-ray absorption becomes detectable by CXO only for energies less than 0.4 keV (Chu et al. 2001).

4. DISCUSSION

The comparisons of X-ray and extinction images in the preceding section reveal striking correspondences between the spatial distributions of $A_V$ and the X-ray surface brightnesses of BD $+30\degr$3639 and NGC 7027. These results suggest that intranebular extinction plays a very important role in determining the X-ray emission morphologies of young dusty PNe. The close correspondence of regions of low extinction and bright X-ray emission in both BD $+30\degr$3639 and NGC 7027 suggests that some extended X-ray emission may remain undetected, making it difficult to draw conclusions as to the intrinsic shape (e.g., axisymmetric vs. elliptically symmetric) of their soft X-ray-emitting regions.

Extinction alone cannot fully explain the strong departures from spherical symmetry and small-scale clumping of the X-ray emission from all three nebulae, however. If extinction were the sole contributor to the X-ray asymmetries of BD $+30\degr$3639 and NGC 7027, then we would expect these asymmetries to be more extreme at lower energy, at which photons are easily absorbed, than at higher energy, at which the X-rays are more highly penetrating. By way of analogy, we note that the near-infrared emission-line morphologies of both nebulae are substantially more symmetric and regular than their optical emission-line morphologies, a difference that can be ascribed to the large decrease in the probability of scattering or absorption of photons by intranebular dust in the near-infrared, relative to the optical regime.

In the X-ray imagery presented here, however, only NGC 7027 seems to display the expected trend of increasing symmetry with increasing energy—although even at high energy, its X-ray morphology differs sharply from its shell-like appearance in near-infrared emission-line images (Fig. 7). BD $+30\degr$3639 is observed to become clumpier and more one-sided as X-ray energy increases. The hard-band emission from NGC 6543, for which extinction effects are negligible in the bands considered here, is mostly confined to a small region of the north lobe. It seems, therefore, that strongly position-dependent variations in the physical conditions in the shocked gas are responsible, at least in part, for the asymmetric and/or clumpy X-ray-emitting regions of these PNe. Indeed, while it seems likely that intranebular extinction modulates the nebular X-ray surface brightnesses of BD $+30\degr$3639 and NGC 7027, we cannot as yet rule out other interpretations. For example, the regions of brightest X-ray emission may correspond to the smallest $A_V$ because...
the dust in these regions is being destroyed by the intense high-energy radiation.

Soker & Kastner (2003) argue that the X-ray luminosities and temperatures of PNe can be explained by a model in which the diffuse emission is produced by a shocked, moderate-speed (\( \sim 500 \, \text{km s}^{-1} \)) post-AGB wind or a wind emanating from a companion to the central post-AGB star, in which the shocked gas has had time to expand and cool adiabatically. This wind and the distribution of previously ejected AGB material also should largely determine the spatial distribution of the X-ray–emitting gas. The post-AGB or companion wind may be highly collimated, resulting in strong axial or point symmetries observable in X-rays. Other processes might then govern the temperature structure (and therefore spatial distribution) of X-ray–emitting gas in detail, with heat conduction from shock-heated to photionized gas in the presence of magnetic fields and/or direct mixing of shock-heated and photionized gas being the most likely candidates (Soker & Kastner 2003 and references therein). These temperature-modulating mechanisms may be more or less important in different nebulae.

The images presented here, combined with the global X-ray properties of PNe displaying diffuse X-ray emission (Soker & Kastner 2003, their Table 1), support this general model. That is, adiabatic cooling of a shocked, moderate-speed, post-AGB or companion wind should produce X-ray luminosities and temperatures in the general range observed \([\sim10^{32} \, \text{ergs s}^{-1}}\) and \((1-3) \times 10^{6} \, \text{K}, \) respectively. The detailed X-ray surface brightness distributions of NGC 6543 and NGC 7027 suggest, moreover, that collimated outflows are responsible for the detailed structure of their X-ray–emitting regions.

In NGC 7027, which exhibits collimated outflows in near-infrared line emission (Cox et al. 2002), the action of such a highly directed (bipolar) fast wind from the central star—or, more likely, an as-yet undetected companion—would readily explain its axisymmetric appearance in the hard-band image (Fig. 9) and the apparent “breakout” of X-ray–emitting plasma beyond the bright rim of near-infrared emission from ionized gas (Fig. 7). Indeed, the morphology of high-velocity Br\( \gamma \) emission is very similar to that of the X-ray emission (Cox et al. 2002). Similarly, both the optical emission and the diffuse soft X-ray emission from NGC 6543 appear to be point-symmetric (see Fig. 10 and Chu et al. 2001), suggestive of the action of collimated outflows (although Chu et al. ascribe the limb-brightened appearance of the X-ray emission from NGC 6543 to the effects of heat conduction or mixing of shocked plasma with photionized gas).

\( \text{BD} +30^\circ 3639 \) may be an example of a PN in which heat conduction and/or mixing play a crucial role in determining X-ray morphology. Either mechanism would explain the bright compact X-ray–emitting region at the eastern rim of the nebula, with heat conduction more naturally explaining its plasma abundances (Soker & Kastner 2003). However, \( \text{BD} +30^\circ 3639 \) and NGC 7027 bear strong resemblances in many fundamental respects (e.g., Soker & Kastner 2003), and, in particular, \( \text{BD} +30^\circ 3639 \) is also known to harbor collimated outflows (Bachiller et al. 2000). Hence, the asymmetric appearance of BD \(+30^\circ 3639\) at high energy, in comparison with NGC 7027, may be primarily the result of the different line-of-sight inclinations of these two nebulae (Masson 1989; KSVDOO; see, however, Bryce & Mellema 1999, who find that BD \(+30^\circ 3639\) more closely resembles the elliptical PN NGC 40). It remains to demonstrate, however, how collimated outflows might produce the off-center X-ray–bright spot in BD \(+30^\circ 3639\), as the outflows detected by Bachiller et al. do not correspond spatially to this region of X-ray emission.

5. SUMMARY AND CONCLUSIONS

We have presented and analyzed reprocessed X-ray \textit{Chandra X-Ray Observatory} images of the planetary nebulae BD \(+30^\circ 3639\), NGC 7027, and NGC 6543. To maximize the spatial resolution of the images, we employ an event relocation technique that takes full advantage of \textit{CXO}'s exceptional spatial resolution; we estimate that the reconstructed images represent a \( \sim 30\% \) improvement in PSF FWHM over the original images. Energy-resolved images demonstrate that the morphologies of all three nebulae depend sensitively on energy. Whereas NGC 7027 appears clumpiest at low energy and more symmetric at higher energy, BD \(+30^\circ 3639\) becomes more one-sided and asymmetric with increasing X-ray energy. Diffuse emission from NGC 6543 is dominated by soft X-rays, with its central point source dominant in the harder energy bands.

With the aid of high-resolution optical and near-infrared H\( \alpha \) emission-line images, we consider the extent to which nonuniform intranuclear absorption affects the X-ray morphologies of these PNe. For both BD \(+30^\circ 3639\) and NGC 7027, we find a strong anticorrelation between broadband X-ray surface brightness and the degree of visual extinction, where the latter is inferred from the spatial distribution of H\( \alpha \) line ratios. We conclude that extinction plays an important role in the overall X-ray appearances of these young, dusty, molecule-rich PNe. The PN NGC 6543 has a much smaller and more spatially uniform degree of visual extinction than either BD \(+30^\circ 3639\) or NGC 7027, however, so spatially variable absorption does not affect its X-ray morphology.

While extinction is important in determining the X-ray surface brightnesses of BD \(+30^\circ 3639\) and NGC 7027, extinction alone cannot explain many important characteristics of the X-ray morphologies of these PNe. In particular, the highly asymmetric appearance of BD \(+30^\circ 3639\) and the axisymmetric or point-symmetric appearances of NGC 7027 and NGC 6543 appear to point out intrinsic X-ray emission properties of these nebulae. We conclude that the action of collimated outflows likely dictates the X-ray appearances of at least two of these nebulae studied here, with heat conduction in the presence of magnetic fields potentially playing a more important role in the asymmetric surface brightness of BD \(+30^\circ 3639\).

Additional spatially resolved \textit{CXO}/ACIS and \textit{XMM-Newton} observations of planetaries are required to better establish the relative importance of the various processes—e.g., collimated outflows, heat conduction, mixing of shock-heated and photionized plasma—proposed to explain their X-ray morphologies as well as their X-ray luminosities and

\footnote{It seems unlikely that a moderate-speed (\( \sim 400 \, \text{km s}^{-1} \)) collimated wind could be driven by the present central star, as such a wind speed is much smaller than the escape velocity from a star of mass of 0.7 \( M_\odot \) and radius 0.07 \( R_\odot \). (Latter et al. 2000b).}
temperatures. We have demonstrated here that these observations must be accompanied by subarcsecond emission-line imaging in the optical and near-infrared, to unambiguously distinguish between intrinsic and extrinsic processes that can determine the energy-dependent X-ray morphologies of PNe.

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